

This item is the archived preprint of:

An RR Lyrae family portrait : 33 stars observed in Pisces with K2-E2

Reference:

Molnar L., Szabo R., Moskalik P.A., Nemeč J.M., Guggenberger E., Smolec R., Poleski R., Plachy E., Kolenberg Katrien, Kollath Z.- An RR Lyrae family portrait : 33 stars observed in Pisces with K2-E2

Monthly notices of the Royal Astronomical Society / Royal Astronomical Society [London] - ISSN 0035-8711 - 452:4(2015), p. 4283-4296

Full text (Publishers DOI): <http://dx.doi.org/doi:10.1093/mnras/stv1638>

Handle/Permalink: <http://hdl.handle.net/10067/1278310151162165141>

An RR Lyrae family portrait: 33 stars observed in Pisces with K2-E2

L. Molnár^{1,2*}, R. Szabó¹, P. A. Moskalik³, J. M. Nemeč⁴, E. Guggenberger^{5,6},
R. Smolec³, R. Poleski⁷, E. Plachy^{1,2}, K. Kolenberg^{8,9,10}, Z. Kolláth^{1,2}

¹*Konkoly Observatory, MTA CSFK, Konkoly Thege Mikós út 15-17, H-1121 Budapest, Hungary*

²*Institute of Mathematics and Physics, Savaria Campus, University of West Hungary*

Károlyi Gáspár tér 4, H-9700 Szombathely, Hungary

³*Copernicus Astronomical Center, Polish Academy of Sciences, ul. Bartycka 18, 00-716 Warszawa, Poland*

⁴*Department of Physics & Astronomy, Camosun College, Victoria, British Columbia, V8P 5J2, Canada*

⁵*Max Planck Institut für Sonnensystemforschung, Justus-von-Liebig-Weg 3, 37077 Göttingen, Germany*

⁶*Stellar Astrophysics Centre, Department of Physics and Astronomy, Aarhus University, DK-8000 Aarhus C, Denmark*

⁷*Department of Astronomy, Ohio State University, 140 W. 18th Ave., Columbus, OH 43210, USA*

⁸*Instituut voor Sterrenkunde, Celestijnenlaan 200D, B-3001 Leuven, Belgium*

⁹*Harvard-Smithsonian Center for Astrophysics, 60, Garden street, Cambridge MA 02138, USA*

¹⁰*Department of Physics, University of Antwerp, Groenenborgerlaan 171, 2020 Antwerp, Belgium*

Accepted 1988 December 15. Received 1988 December 14; in original form 1988 October 11

ABSTRACT

A detailed analysis is presented of 33 RR Lyrae stars in Pisces observed with the *Kepler* space telescope over the 8.9-day long K2 Two-Wheel Concept Engineering Test. The sample includes not only fundamental-mode and first overtone (RRab and RRc) stars but the first two double-mode (RRd) stars that *Kepler* detected and the only modulated first-overtone star ever observed from space so far. The precision of the extracted K2 light curves made it possible to detect low-amplitude additional modes in all subtypes. All RRd and non-modulated RRc stars show the additional mode at $P_X/P_1 \sim 0.61$ that was detected in previous space-based photometric measurements. A periodicity longer than the fundamental mode was tentatively identified in one RRab star that might belong to a gravity mode. We determined the photometric [Fe/H] values for all fundamental-mode stars and provide the preliminary results of our efforts to fit the double-mode stars with non-linear hydrodynamic pulsation models. The results from this short test run indicate that the K2 mission will be, and has started to be, an ideal tool to expand our knowledge about RR Lyrae stars. As a by-product of the target search and analysis, we identified 165 *bona-fide* double-mode RR Lyrae stars from the Catalina Sky Survey observations throughout the sky, 130 of which are new discoveries.

Key words: stars: variables: RR Lyrae

1 INTRODUCTION

RR Lyrae stars are pulsating low-mass horizontal-branch stars. Since they can be found at the intersection of the horizontal branch and the classical instability strip, a relatively small region of the Hertzsprung-Russell-diagram, their absolute brightnesses are fairly similar, making them good distance indicators. The large number of RR Lyrae stars make them good tracers of halo structures around the Milky Way (Sesar et al. 2013) and even in nearby

galaxies (e.g. Brown et al. 2004; Sarajedini et al. 2006; Haschke, Grebel & Duffau 2012). The period ratios of the rare double-mode stars can be used to determine the masses of RR Lyrae stars independent from stellar evolution theory.

The *Kepler* space telescope observed ~ 45 RR Lyrae stars in the Lyra-Cygnus field of the prime mission, spanning a wide metallicity and brightness range, from the eponym of the class, RR Lyr ($Kp = 7.862$ mag, Kolenberg et al. 2011), down to about $Kp = 17.5$ magnitudes (Nemeč et al. 2013). Although the *Kepler* mission was designed to detect subtle light changes caused by planetary transits (Borucki et al. 2010) and the high-amplitude variations of the RR Lyrae

* E-mail: molnar.laszlo@csfk.mta.hu

stars required additional processing, space-based photometry proved to be essential to further our understanding of RR Lyrae stars (Benkő et al. 2014).

Kepler observed mostly fundamental-mode stars, with only four first-overtone and no double-mode stars in the sample. Therefore most of its findings have been limited to the RRab class. The results include the discovery of period doubling (Kolenberg et al. 2010; Szabó et al. 2010), a wealth of low-amplitude additional modes (Benkő et al. 2010; Guggenberger et al. 2012) and the discovery that most Blazhko stars show multiple or irregular modulations (Benkő et al. 2014). The observations also revealed that the new phenomena (period doubling and additional modes) are strictly limited to the modulated stars (Nemec et al. 2011), and a recent analysis of stars observed by the *CoRoT* space telescope confirmed this dichotomy (Szabó et al. 2014). In contrast, all four first-overtone stars observed by *Kepler* (and two more in the *CoRoT* sample) exhibit the same additional mode at the frequency ratio $f_1/f_2 \approx 0.60 - 0.64$, without any signs of modulation in the main mode (Moskalik et al. 2013, 2015). The origin of this additional mode is uncertain. We note that we refer to this mode as f_X instead of f_2 to avoid any confusion with the suspected second-overtone signals in the RRab stars.

The primary mission of *Kepler* gathered four years of (quasi)-continuous observations. These allowed for more specific studies, such as following the very unexpected vanishing of the Blazhko effect in RR Lyr (Le Borgne et al. 2014) or the first attempt to search for chaos in the modulation of an RRab star (Plachy et al. 2014), and even the search for artificial origins in the period variations (Hippke et al. 2014). After the unfortunate failure of the second reaction wheel, however, maintaining the original pointing was not feasible any more.

The new K2 mission of the *Kepler* space telescope uses orientations along the Ecliptic where the spacecraft can balance against the radiation pressure of the Sun (Howell et al. 2014). This new strategy results in short, ~ 75 day long observing campaigns that scan various fields in the sky, reaching targets the original field lacked. In this paper we detail the results from the very first measurements of the K2 mission. Section 2 describes the K2-E2 run and the extraction of the light curves. Results are detailed in Sections 3 and 4 for RRd-RRc and RRab stars, respectively. Conclusions and future directions are summarized in Section 5.

2 K2 DATA AND REDUCTION

Kepler observed a stellar field around the vernal equinox point in Pisces (center coordinates: $\alpha = 359^\circ$, $\delta = -2^\circ$) between 4th-13th February 2014. The primary goal of this K2 Two-Wheel Concept Engineering Test (hereafter K2-E2) was to test the performance of the telescope in fine guidance mode. As well, the observations of nearly 2000 targets were made available for the scientific community. Instead of the tight target pixel masks used in the primary mission, larger subframes were stored around the targets in order to avoid any flux loss due to the drift of the telescope. A ~ 1 pixel variation in the position of the stars was observed during the 8.9-day long K2-E2 data. The boresight was shifted by multiple pixels 2.3 days into the test, at BJD = 56695.36, when

Table 1. Sample of the complete K2-E2 light curve file. The full table is accessible online.

EPIC type: RRd	BJD-2450000 (d)	K_p (mag)	ΔK_p (mag)
60018653	1860.0503350	13.90840	0.00054
60018653	1860.0707686	13.92217	0.00054
60018653	1860.0912023	13.92438	0.00054
60018653	1860.1116358	13.91703	0.00054
60018653	1860.1320694	13.89827	0.00053
...			

the telescope locked successfully into fine guidance mode. Fortunately the 50x50-pixel target masks were more than enough to contain these movements.

We identified 34 potential RR Lyrae stars in the K2-E2 sample and extracted their photometric data with the PYKE software¹, developed for the *Kepler* mission by the *Kepler* Guest Observer Office (Still & Barclay 2012). All targets were observed in long cadence mode, with a sampling of 29.4 minutes. To counter the effect of drifting and capture all the flux, we defined relatively large pixel apertures for the stars, with 1–2-pixel-wide halos around the actual point spread functions. We applied two different masks for the positions before and after the boresight shift. We used the background-corrected SAP (Simple Aperture Photometry) fluxes calculated by the KEPEXTRACT routine. The photometric precision spreads between 3.5 μ mag to 7 mmag per data point for a 10.5 and a 17.4 mag star, respectively. In some cases we had to apply small amounts of shift and/or scaling to the data to connect the two parts of the light curves. These differences were likely caused by different pixel sensitivities in the two positions. Some stars are affected by the Fine Guidance Sensor clock crosstalk patterns (van Cleve & Caldwell 2009). The faintest star in the sample, EPIC 60018780, suffers from extensive negative video crosstalk from a bright target elsewhere on the same module: dark CCD columns cross the image of the star. As the image of the star moves in and out of the dark column, jumps up to 0.1–0.2 mag appear in the light curve of this star, especially in the first 2.3 days.

Without the third reaction wheel, radiation pressure from the Sun slowly turns *Kepler* away from the desired spacecraft attitude, and rolls it about the optical axis (the X-axis of the spacecraft, Howell et al. 2014). These disturbances are periodically corrected with the on-board thrusters. For some stars the frequency of the manoeuvres, $f_{\text{corr}} = 8.16 \text{ d}^{-1}$, and its harmonics can be detected in the Fourier spectra. (We note that the K2-E2 run tested a 3-hour adjustment frequency, but later campaigns use 6-hour intervals, corresponding to $f_{\text{corr}} = 4.08 \text{ d}^{-1}$.) Rolling of the field-of-view means that stars close to the center are least affected but targets close to the edges can experience considerable drifting. Although these peaks are well-defined and can be easily excluded from the analysis, they indicate that special care is needed to define the target apertures and to correct for pixel sensitivity variations.

Undoubtedly, there will be better and more precise

¹ <http://keplergo.arc.nasa.gov/PyKE.shtm1>

Table 2. Sample of the frequency tables of the entire K2-E2 data set: frequencies, amplitudes, phases, the respective uncertainties and the identifications of the peaks (f_0 : fundamental mode, f_1 : first overtone, f_X , f_2 , f_g : additional modes). Super-Nyquist frequency components are marked with asterisks. The entire table is accessible in the online version of the paper.

Star	F (d $^{-1}$)	A (Kp mag)	ϕ (rad/2 π)	ΔF (d $^{-1}$)	ΔA (Kp mag)	$\Delta\phi$ (rad/2 π)	ID
60018653	1.85382	0.04641	0.92372	0.00031	0.00040	0.00080	f_0
60018653	2.48564	0.13988	0.82381	0.00010	0.00035	0.00026	f_1
60018653	3.7076	0.00422	0.2164	0.0034	0.00034	0.0088	$2f_0$
60018653	5.561	0.00105	0.573	0.014	0.00037	0.035	$3f_0$
60018653	4.97128	0.02273	0.1761	0.00063	0.00036	0.0016	$2f_1$
...							

ways to extract K2 photometry as the mission evolves. However, (semi-)automated methods are usually optimised for relatively quiet stars and not for large-amplitude, short-period pulsators. One such pipeline, developed by Vanderburg & Johnson (2014), for example, admittedly failed to improve the light curves of RR Lyrae stars and eclipsing binaries. Therefore we were satisfied with the level of accuracy we achieved with the above method for this initial study. The light curves (along with the frequency tables, see Section 2.1) are available online for further analysis alongside the paper and at the *Kepler* Investigations at Konkoly webpage². Samples are provided in Tables 1 and 2.

2.1 The K2-E2 RR Lyrae sample

Prior to the K2-E2 test, a call was issued within KASC (the *Kepler* Asteroseismic Science Consortium) to provide targets within the approximate position of the field in Pisces. We proposed 143 RR Lyrae stars, largely based on the RRab catalogue of the Catalina Sky Survey (CSS, Drake et al. 2013a,b) but also on the ASAS database and the SIMBAD identifications. By the time the K2-E2 data became available, the detailed classifications of periodic variables in CSS were published (Drake et al. 2014), allowing for a follow-up search in the target masks. We found 34 stars at or close to the positions of the 1952 K2-E2 targets based on these catalogues. Comparison with the proposed list revealed that several of the originally proposed targets ended up off silicon, while others were not selected for observation.

One of the stars, EPIC 60018241 (ASAS J234439-0148.6) turned out to be an eclipsing binary (Conroy et al. 2014), but the other 33 stars are RR Lyrae stars. Surprisingly, we identified RR Lyrae stars from all subclasses, including two classical double-mode (RRd) stars and a potential strongly modulated first-overtone (RRc) candidate. Basic parameters, i.e. the EPIC (*Kepler* Ecliptic Plane Input Catalog) identification number, coordinates, brightness, period, other identifications, and subclass, are summarised in Table 3 for RRd and RRc stars. The parameters of RRab stars are listed in 5.

We used the `Period04` software to calculate the Fourier transforms (Lenz & Breger 2005). The frequency tables of the K2-E2 sample are summarised in Table 2. Here we included only the resolved frequency peaks, i.e. unresolved residuals that indicate amplitude and/or phase variation are not listed. In some cases high-order harmonics reflect back

from the Nyquist frequency ($f_N = 24.4684$ d $^{-1}$), but we fitted the super-Nyquist values in all cases. These are marked with asterisks in the frequency table.

3 DOUBLE-MODE AND FIRST-OVERTONE STARS

3.1 Double-mode stars

We identified two RRd-type stars in the field. Both stars, EPIC 60018653 (CSS J235643.2+030119) and 60018662 (CSS J234207.3+003253), were misclassified in the Catalina sample, as RRc and RRab stars, respectively, but the K2 light curves revealed that they are actually classical double-mode RR Lyraes (Figure 1). These are the first RRd stars observed by *Kepler* and only the third and fourth from space. AQ Leo, the brightest of the class, was observed by *MOST* (Gruberbauer et al. 2007), and the *CoRoT* data of another star, CoRoT ID 0101368812, was analysed by Chadid (2012) and Szabó et al. (2014).

3.1.1 Period ratios and physical properties

The two K2 stars have similar pulsation properties. Their fundamental-mode periods are $P_0 = 0.53943$ d and $P_0 = 0.55900$ d, while the period ratios are $P_1/P_0 = 0.74581$ and $P_1/P_0 = 0.74678$ for EPIC 60018653 and 60018662, respectively. These values put them into the long-period end of the Petersen diagram that displays the period ratios against the fundamental-mode period of the stars (Figure 2). In fact, EPIC 60018662 is near to the long-period edge of the area covered by double-mode stars.

We constructed a Petersen diagram to put the two stars into broader context. The diagram in Figure 2 includes the largest sample of field RRd stars constructed so far. One of us (RP) searched for bona-fide double-mode stars in the Catalina Surveys Catalog of Periodic Variable Stars (Drake et al. 2014) and identified 130 new RRd variables. The details of this search and the list of stars are presented in Appendix A. We also included the stars identified in the LINEAR (Poleski 2014), ASAS/NSVS (Wils 2005; Wils, Lloyd & Bernhard 2006; Bernard & Wils 2006; Szczygiel & Fabrycky 2007), and SuperWASP (Wils 2010) databases. Stars from the Sagittarius dwarf galaxy and the OGLE bulge survey provide the background (grey crosses; Cseresnjés 2001; Soszyński et al. 2014). It is interesting to note that even with the CSS sample, field stars outside the bulge seem to be restricted to $P_0 \geq 0.45$ d and $P_1/P_0 \geq$

² <http://konkoly.hu/KIK/data.html>

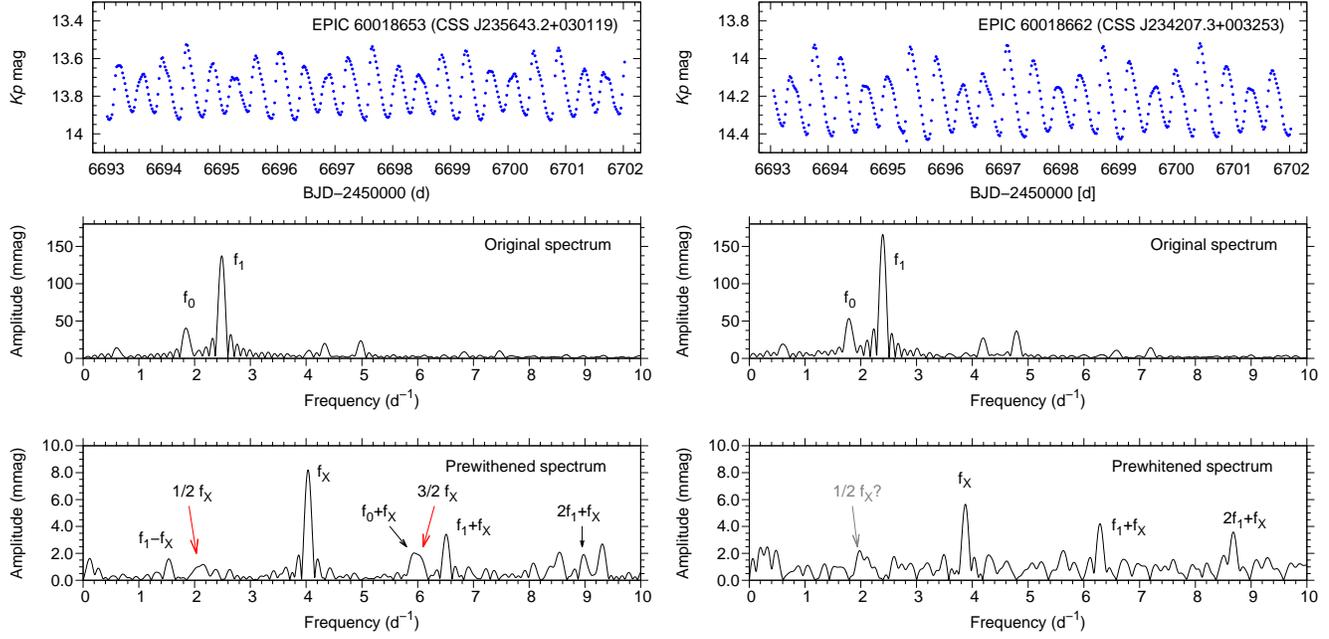


Figure 1. Light curves (top row) and Fourier-spectra of the two RRd stars in the sample. Middle row: the original Fourier spectra. Fundamental-mode and first-overtone frequency peaks are labelled with f_0 and f_1 . Bottom row: residual spectra after prewhitening the data with f_0 , f_1 and their linear combination peaks. The additional mode f_X and its significant subharmonic peaks and combinations are labelled. We also indicated the marginally detected $1/2 f_X$ peak in EPIC 60018662.

Table 3. Double-mode and first-overtone RR Lyrae stars observed during the Two-Wheel Concept Engineering Test run of the K2 mission. O1 and FM indicate the periods of the first overtone and fundamental mode, respectively. We provide estimated Kp magnitudes for two stars (in italics). The brightness of EPIC 60018238 is too high in the EPIC ($Kp = 9.83$), compared to the flux values. The last star has no EPIC ID but was observed within the pixel mask of the engineering target EPIC 60042292. **Modulation period.

EPIC ID	RA (deg)	Dec (deg)	Kp mag	O1 period (d)	Name	FM period (d)	CSS ID number
60018653	359.18015	+3.022110	13.760	0.402311	CSS J235643.2+030119	0.539427	1104128004573
60018662	355.53058	+0.547972	14.216	0.417448	CSS J234207.3+003253	0.558996	1101127011298
60018224	358.53385	+0.963635	10.600	0.306156	EV Psc	–	–
60018238	354.87500	−3.631667	<i>12.419</i>	0.274804	ASAS J233930−0337.9	–	–
60018678	352.16375	+1.028500	14.862	0.432426	CSS J232839.2+010142	–	1101126021018
(60042292)	359.42500	−1.839444	<i>14.502</i>	0.300110	CSS J235742.1−015022	17.07**	1001128017680

0.742, whereas the Sgr and bulge population extends down to $P_0 \approx 0.35$ d and $P_1/P_0 \approx 0.726$ (Soszyński et al. 2011).

The precise period values of pulsation modes in a star depend strongly on the fundamental parameters: mass, luminosity, effective temperature, and metallicity of the star. Therefore the simultaneous detection of two radial modes in one target can put strong constraints on its physical parameters. This is especially important for mass determination, as we currently have no other means to measure RR Lyrae masses. Although a few promising binary candidates were recently discovered by Hajdu et al. (2015), dynamical masses will take several years if not decades to measure. According to the calculations of Szabó et al. (2004), RRd stars are confined to a narrow parameter range. We carried out a pilot study to see whether we can find good hydrodynamic models effectively.

First we estimated the effective temperature, mass, luminosity and metal content ranges based on the pulsation

model results of Szabó et al. (2004). Compared to their Figure 6/a, the period parameters of the two stars fall well below the $Z = 0.001$ models, therefore we selected two OPAL opacity tables with $Z = 0.0001$ and 0.0003 metal contents.

Then a grid of non-linear hydrodynamic models was computed with the Florida-Budapest turbulent convective code (details of the code are presented by Kolláth & Buchler (2001) and Kolláth et al. (2002)). Luminosity was varied between $L = 48 - 60 L_\odot$, and two temperature and mass ranges were selected: one between $T_{eff} = 6480 - 6540$ K, $M = 0.72 - 0.80 M_\odot$, and the other between $T_{eff} = 6410 - 6470$ K, $M = 0.76 - 0.84 M_\odot$. Increments were 2 or 3 L_\odot , 10 K and 0.2 M_\odot .

Models were iterated for 2000 fundamental-mode cycles and then the periods of both modes were calculated. We hand-picked 8 models close to the periods of EPIC 60018653 and 6 models close to 60018662 and iterated them for a further 2000 cycles. After that, 7 models converged slowly

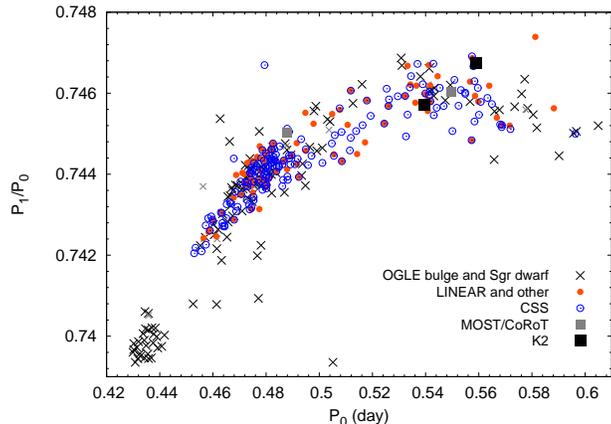


Figure 2. Period ratio versus the period of the fundamental mode of the RRd stars, also known as the Petersen diagram. Small grey crosses: OGLE bulge and Sgr dwarf stars, their sequence continues to even smaller values. Black squares: the two K2 stars; orange dots: stars detected in the LINEAR, ASAS, NSVS and SuperWASP surveys; blue circles: our CSS sample from Appendix A; grey squares: CoRoT ID 0101368812 (left) and AQ Leo (right). Multiple identifications were not filtered. References for the data sources can be found in Section 3.1.1.

to single-mode pulsation and only 7 double-mode models remained. Six of those models are actually pairs that differ only in the temperature step. The results agree with the findings of Szabó et al. (2004) that non-linear RRd models exist only in narrow parameter ranges.

As Figure 3 illustrates we could not fit both stars with the same metal content value. For EPIC 60018653, two models at $T_{eff} = 6410\text{--}20$ K, $M = 0.76 M_{\odot}$, $L = 48 L_{\odot}$ and $Z = 0.0003$ are relatively close to the star on the Petersen diagram. Two models fall very close to EPIC 60018662 in this parameter plane: $T_{eff} = 6430\text{--}40$ K, $M = 0.78 M_{\odot}$, $L = 52 L_{\odot}$ and $Z = 0.0001$. We included, for comparison purposes, three more models that have similar fundamental-mode periods but different period ratios than the stars. Details of all seven models are included in Table 4.

The mass values appear to be quite high: pulsation masses for RRab stars in the original *Kepler* field, for example, fell below $0.66 M_{\odot}$ (Nemec et al. 2011). Unfortunately, in the absence of proper dynamic masses, all mass estimates are model-dependent. Low-mass linear model calculations can be fitted to the whole range of the Petersen diagram (Soszyński et al. 2011). However, non-linear RRd models in the Florida-Budapest code appear only at higher masses, typically above $0.7 M_{\odot}$ and they are still consistent with evolutionary models. We also note that the ability of current non-linear pulsation codes to properly model double-mode pulsation is still a matter of debate; see, e.g. Kolláth et al. (2002) and Smolec & Moskalik (2008) for opposing views. We also note that the lack of the f_X mode in the models does not invalidate our findings. If we assume that it is a non-radial mode, it could naturally arise in higher-dimension models with same the fundamental parameters.

We emphasize that these are just preliminary results from a pilot study. Calculating double-mode non-linear models is time-consuming, mainly because mode amplitudes may change very slowly. Nevertheless, we plan to carry out a more

Table 4. Hydrodynamic model results for the two RRd stars, numbers correspond to EPIC 60018653 and 60018662. We also included the models where the first-overtone periods do not fit the observations, as indicated in Figure 3.

Star	Mass (M_{\odot})	Lum. (L_{\odot})	T_{eff} (K)	Z	P_0 (d)	P_1/P_0
Good model fits						
653	0.76	48	6410	0.0003	0.54194	0.74538
653	0.76	48	6420	0.0003	0.53906	0.74545
662	0.78	52	6430	0.0001	0.56036	0.74665
662	0.78	52	6440	0.0001	0.55737	0.74674
Models with different period ratios						
(653)	0.76	48	6410	0.0001	0.53917	0.74689
(662)	0.78	51	6410	0.0003	0.56031	0.74516
(662)	0.78	51	6420	0.0003	0.55731	0.74524

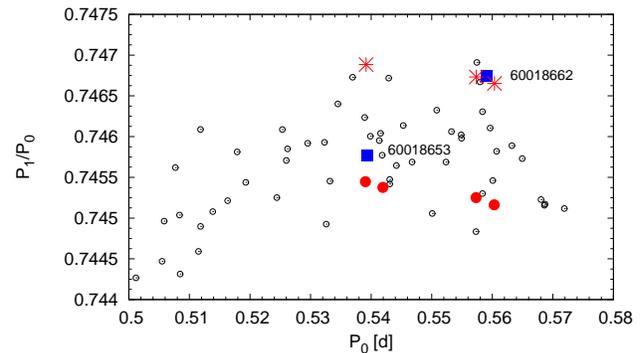


Figure 3. Hydrodynamic model results compared to the two RRd stars (blue squares) on the Petersen diagram. The small circles are the CSS RRd stars from Appendix A. Large red dots are the $Z = 0.0003$ models, large crosses are the $Z = 0.0001$ models.

extensive survey after more RRd light curves will become available from the K2 mission.

3.1.2 Additional modes in RRd stars

The stars AQ Leo and CoRoT ID 0101368812 both showed additional modes that could not be explained by linear combinations of the fundamental mode and the first overtone. Comparison with the Fourier-parameters we derived from the K2 stars revealed striking similarities to these double-mode stars. One particular mode that falls to the frequency ratio of $f_1/f_X \approx 0.615$, was identified in both stars, and it has been recently identified in a modulated RRd star from the OGLE database (Smolec et al. 2015) as well. Figure 1 shows that we detected the same additional mode at frequency ratios $f_1/f_X = 0.6162$ in EPIC 60018653 and $f_1/f_X = 0.6166$ in 60018662. The same mode was reported not only in RRd stars but in RRc stars as well (Section 3.2).

We detected the subharmonics, or half-integer peaks of the f_X mode ($1/2$ and $3/2 f_X$) in EPIC 60018653. The presence of these frequencies indicates that the mode experiences period doubling. The peaks are broadened and their centres are slightly off from the exact half-integer values at $0.499 f_X$ and $1.495 f_X$, respectively. Temporal vari-

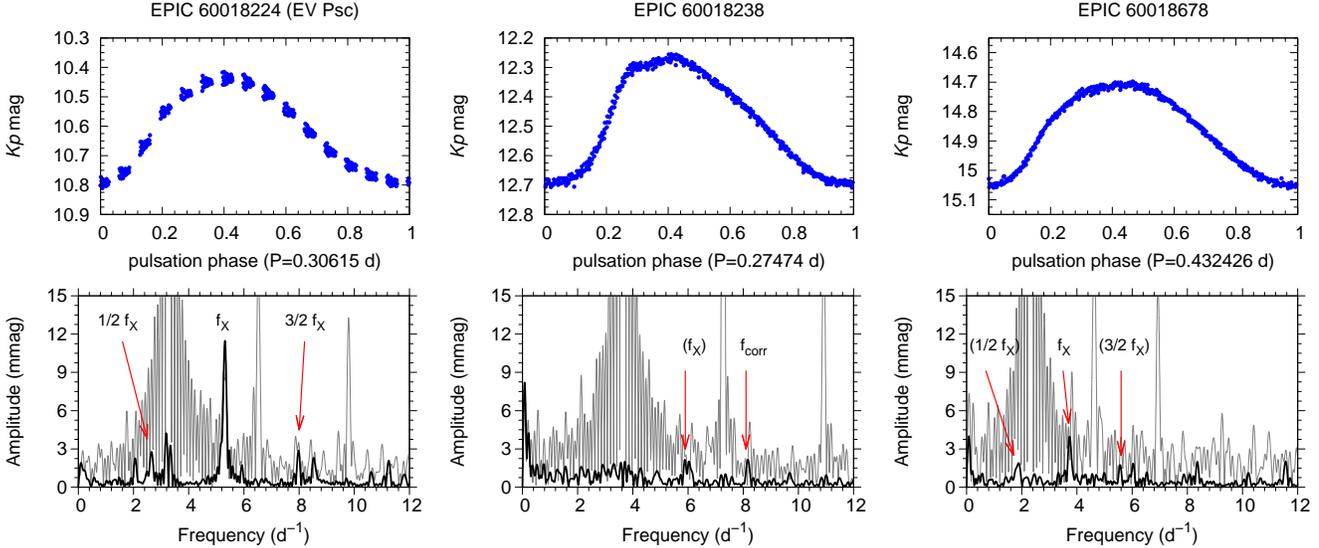


Figure 4. Folded light curves and Fourier-spectra of the three non-modulated RRc stars. Original spectra are shown with thin grey lines; the black thick lines are the residuals after prewhitening with the $n f_1$ ($n = 1, 2, \dots, 5$) frequencies. The f_X peak and the corresponding half-integer peaks are indicated. f_{corr} is the instrumental signal of the correction manoeuvres. Labels of low-significance peaks (S/N ratio between 3 and 4) are in parentheses.

ations in period doubling created forests of peaks in the spectra of RRab stars observed in the nominal mission and the highest-amplitude peaks were also displaced from the exact values (Szabó et al. 2010). Temporal variations have been detected in the f_X mode and its subharmonics in the *Kepler* RRc sample (Moskalik et al. 2015). It is very likely that the broadened peaks in the K2-E2 sample are caused by temporal variations in the period doubling of the f_X mode as well. The $1/2 f_X$ peak is marginally detectable in EPIC 60018662 too, but only with a signal-to-noise ratio of 2.5.

Gruberbauer et al. (2007) observed the same $1/2 f_X$ peak in AQ Leo, but identified this peak tentatively as the parent frequency, f_i , and the true f_X additional mode as $2 f_i = f_{ii}$. Chadid (2012) reported two other additional modes (f_3, f_4) in the *CoRoT* star. However, the frequency of their $f_4 - f_1$ combination peak is equal to $1.49 f_X$. If we accept this peak as a subharmonic, then f_4 itself becomes a combination peak ($1/2 f_X - f_1$) instead of an independent mode, making the star more similar to the rest of the RRc sample. Similar combinations can be detected in EPIC 60018653 too. We can conclude that not only the f_X additional mode, but its period doubling effect is observable in all four RRd stars measured by space photometric missions so far.

3.2 First-overtone stars

Four RRc stars were identified in the K2-E2 data, the same number as observed in the original *Kepler* field. Three of those stars are normal first-overtone pulsators. The fourth star, CSS J235742.1-015022 is very different from the others: it exhibits strong amplitude and phase variations.

The three non-modulated stars have different pulsation periods and light-curve shapes, as illustrated in Figure 4. EPIC 60018238 has the shortest period ($P = 0.27474$ d)

and a very pronounced hump feature just before maximum light. We detected two different additional modes with very small amplitudes. The star seems to exhibit the ubiquitous f_X mode, although we detected two, closely-spaced peaks at $f_1/f_X = 0.6154$ and $f_1/f'_X = 0.6048$ instead of a single frequency, with low significances (S/N ratios are ~ 3.3). We also identified the drift-correction frequency at 8.16 d^{-1} in the star.

The light curves of the other two RRc stars, EV Psc (EPIC 60018224) and EPIC 60018678 show almost no shock-related features (humps or bumps). The pulsation period of EV Psc exhibits strong beating with the 1766-second long-cadence sampling period, so it is not covered continuously (Figure 4). We detected the f_X mode in both stars at frequency ratios of $f_1/f_X = 0.6146$ (EV Psc) and 0.620 (60018678). The additional mode is very pronounced in EV Psc: the scatter of the folded light curve in Figure 4 is caused primarily by the mode itself. Moreover, we also detected the half-integer frequencies ($1/2 f_X$ and $3/2 f_X$) in EV Psc, corresponding to period doubling in the f_X mode. Half-integer peaks can be located but are marginal in EPIC 60018678. It seems that the (period-doubled) f_X mode is intimately connected to the first radial overtone: it was observed not only in the K2 RRc stars but in the original *Kepler* sample, in two *CoRoT* stars, and the RRd stars described above as well (Moskalik et al. 2015; Szabó et al. 2014).

In fact the f_X mode has been detected in first-overtone Cepheids first, from the data collected by the OGLE surveys (Moskalik & Kołaczowski 2008; Moskalik et al. 2013). More recently, the f_X mode has been identified in 147 RRc stars in the OGLE bulge sample and in 14 RRc and 4 RRd stars in the globular cluster Messier 3, respectively, further confirming its ubiquity (Netzel, Smolec & Moskalik 2015; Jurcsik et al. 2015). Yet the origin of the mode is still unclear: Dziembowski (2012) hypothesised that high-degree

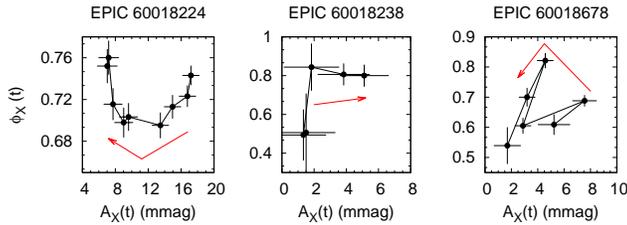


Figure 5. Time variation of the amplitude and phase of the f_X in the non-modulated RRc stars (also known as loop diagrams). Arrows denote the direction of progression. Variation can be detected in all three stars: EPIC 60018224 (EV Psc) shows particularly strong decrease in the amplitude of f_X .

f -modes ($\ell = 42 - 50$) could explain the period ratios but those would require quite high intrinsic mode amplitudes.

3.2.1 Variation of the additional modes

Beside the double peak in EPIC 60018238, we observed residual power in the Fourier spectra of both EPIC 60018224 and 60018678 after prewhitening with f_X . These findings indicate that the f_X mode undergoes amplitude and/or phase variations, similar to the findings of Moskalik et al. (2015) and Szabó et al. (2014) for *Kepler* and *CoRoT* stars, respectively. We split the data sets into 1.5-day (EPIC 60018238 and 60018678) and 1-day (EV Psc) bins and calculated the Fourier-parameters in the bins. As Figure 5 illustrates, all stars show some variation, even on this short time scale. It is of course possible that there are actually multiple unresolved, separate peaks in the spectra but the considerably longer observations in the prime mission showed no signs of closely-spaced peaks in other RRc stars.

We carried out the same analysis on the RRd stars as well, but we did not detect significant amplitude variations in the f_X mode. However, we found the uncertainties of the amplitude and phase data to be much higher than in the RRc stars, possibly because of the more complicated frequency content. Variation in the f_X mode of the star CoRoT ID 0101368812 was detected by Szabó et al. (2014), therefore we expect that the science data of K2 will be sufficient to probe the variation of these modes in other RRc and RRd stars.

3.3 A strongly modulated RRc star

The fourth star, CSS J235742.1-015022 is located within the subframe of an engineering target star, EPIC 60042292, but it has no EPIC identification itself. The K2 observations revealed that it shows very strong amplitude and phase modulation, as seen in Figure 6. The Blazhko effect exists in first-overtone RR Lyrae stars but it is less common than in their fundamental-mode siblings. Mizerski (2003) found a 10 per cent incidence rate in the OGLE-II bulge data while Nagy & Kovács (2006) derived 7.5 per cent for the Large Magellanic Cloud based on MACHO data, although both numbers should be considered as lower limits only. Recent surveys, such as the Northern Sky Variability Survey and the All-Sky Automated Survey identified several modulated field RRc stars (Wils et al. 2006; Szczygiel & Fabrycky

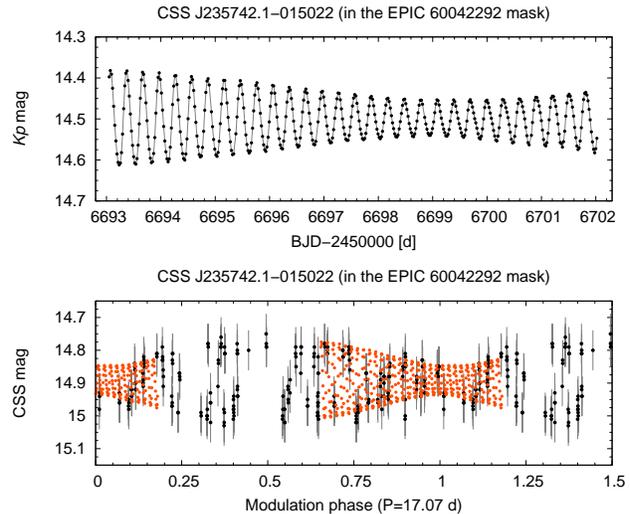


Figure 6. Top: K2-E2 light curve of the modulated RRc star. Bottom: the Catalina Sky Survey (black) and K2 data (orange, shifted to the CSS brightness level), and folded with the modulation period.

2007). The Blazhko Stars in the Galactic Field database currently lists 55 modulated RRc stars (Skarka 2013). First-overtone stars with obvious Blazhko effect have not been observed before from space, though we note that Szabó et al. (2014) found a very small side peak in one of the *CoRoT* RRc stars (ID 0105036241) that might be connected to amplitude and/or phase variations.

3.3.1 Modulation properties of the RRc star

The K2-E2 data covers less than a modulation cycle, therefore we analysed the CSS observations to determine the modulation period. We identified a single potential side peak, $f_1 + f_m$, and determined a tentative modulation period of $P_m = 17.07$ d. Comparison with the K2-E2 data revealed that this is indeed the correct value and it has been stable for more than 8 years, over the duration between the start of CSS and the K2 observations (lower panel of Figure 6). The detection of one modulated star out of eight RRc targets in the current *Kepler* and K2 sample, agrees with the findings of previous studies and confirms the lower occurrence of the Blazhko effect among first-overtone stars. Adding the two *CoRoT* stars, the ratio of modulated stars falls somewhere between 10–20 per cent, depending on whether we consider CoRoT ID 0105036241 modulated or not. More observations will be required of course to determine the occurrence rate more precisely. K2 is perfectly suited to provide the answer: we expect it to observe approximately 50–100 galactic RRc stars over the course of the mission.

We divided the data of CSS J23542.1-015022 into twelve 0.75-day long segments and calculated the amplitudes and phases of the f_1 frequency peak in each bin to investigate the modulation properties of the star. The results are shown in the lower panels of Figure 7. The pulsation amplitude changes rather symmetrically, but the descending branch of the phase variation is much steeper than the ascending branch. In a phase-amplitude diagram, the modulation

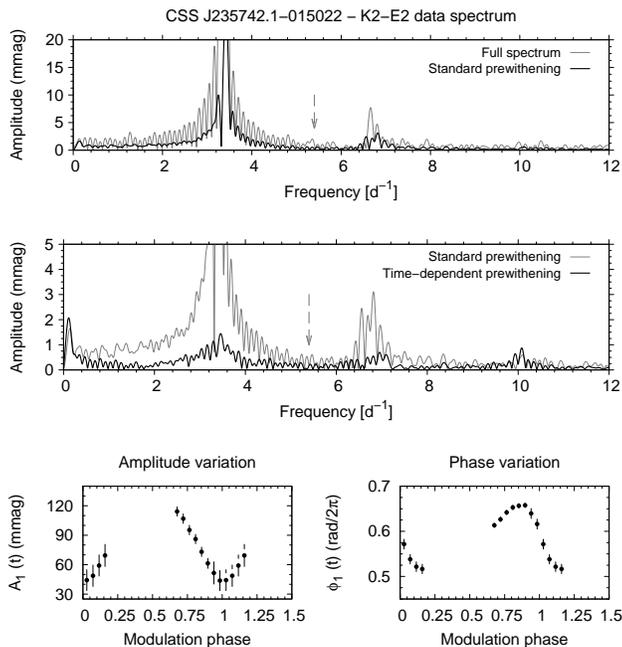


Figure 7. Top: Fourier-spectra of the modulated RRc star, CSS J235742.1-015022 (within the mask of EPIC 60042292), original with thin grey line, prewhitened with thick black line. Middle: residual spectra after standard (thin grey) and time-dependent (thick black) prewhitening with the pulsation frequencies. The dashed arrows marks the expected position of the f_X mode. Bottom: temporal variation of the amplitude and phase of the f_1 frequency, folded with the modulation period.

progresses in the clockwise direction, opposite the direction that was observed in another modulated RRc star, TV Boo (Skarka & Zejda 2013).

3.3.2 Search for additional modes

The Fourier spectrum of the star is shown in the upper and middle panels of Figure 7. We found no signs of additional frequencies in the star. As the K2-E2 data is shorter than the modulation period, the modulation side-peaks cannot be resolved with conventional prewhitening (constant Fourier parameters) properly. We applied an alternative, time-dependent prewhitening method that subtracts non-stationary signals from the data, developed by Moskalik et al. (2015). While this method resulted in much lower residuals, we still could not identify additional modes down to 0.4 millimag amplitude. Interestingly, ground-based observations of two other modulated RRc stars did not reveal any additional modes either, down to about 5 and 2 millimagitudes for LS Her and TV Boo, respectively (Wils, Kleidis & Broens 2008; Skarka & Zejda 2013).

The f_X mode seem to be a common occurrence in normal first-overtone stars that show only low-level temporal variability but not the Blazhko effect (Moskalik et al. 2015). The potential lack of the f_X mode in strongly modulated RRc stars may raise several questions about mode excitation and the physical origin of the modulation. The K2 mission is ideally suited to expand our understanding about

these problems. Further observations can also provide clues to whether the modulation in RRab and RRc stars is the same mechanism or is caused by different physical processes.

4 FUNDAMENTAL-MODE STARS

We found 27 RRab-type stars in the K2-E2 data. Identifications and folded light curves are presented in Table 5 and Figure 8. The brightness range of the stars are similar to the original *Kepler* RR Lyrae sample (excluding RR Lyr itself), between $Kp = 12.199$ and 17.462 magnitudes. We were able to extract high-precision photometry for almost all stars. The image of EPIC 60018780 is affected by dark CCD columns (negative video crosstalk). We extracted the flux from the remaining pixels, but because the pointing drifts affect the dark columns and the star differently, small jumps can be seen in the data, especially in the first part of the observations. Given the limited resolution of *Kepler*, some stars are affected by blending. We see the strongest contaminations from another nearby object in EPIC 60018773 and 60018737. A more detailed evaluation and possibly image subtraction methods will be required to extract better photometry for these stars.

4.1 The Blazhko stars

Although the 8.9-day long time base is rather short, we identified amplitude variations in several stars. 11 out of the 27 stars show significant residual power in their Fourier spectra after the removal of the pulsation frequency and its harmonics. While the data are too short to resolve the modulation side peaks, the light curves show smoothly changing pulsation amplitudes. Therefore it is safe to assume that these stars are modulated. In 13 cases we found no signs of amplitude and/or phase change either in the spectra or in the amplitude variation data. Analysis of 3 stars led to inconclusive results. Therefore at least 40.7 per cent of the RRab stars seem to be modulated, but the number can be as high as ~ 50 per cent, in accordance with previous ground-based and *Kepler* findings (Jurcsik et al. 2009; Nemeč et al. 2013). This is an impressive result if we consider that Blazhko periods are usually much longer than 9 days (in the *Kepler* sample they ranged from 27 days to two years).

In three cases, data problems prevented us from confirming or excluding the presence of modulation. EPIC 60018737 and 60018773 suffer from blending. In both cases, the Fourier spectra did not reveal any clear and significant modulation sidelobes. In fact, slight temporal variation can be detected in the amplitude of the f_1 peak, but the light curve itself is distorted and includes additional variation too. EPIC 60018780 is severely impacted from the negative video crosstalk which resulted in jumps in the light curve and elevated noise level in its Fourier spectrum.

4.1.1 Modulation properties of RRab stars

The K2-E2 data in themselves are too short to determine the modulation periods and amplitudes of the stars. While some of the brighter non-modulated stars were covered by multiple ground-based surveys, for the Blazhko stars we had to rely on the CSS data alone. The Catalina data has very

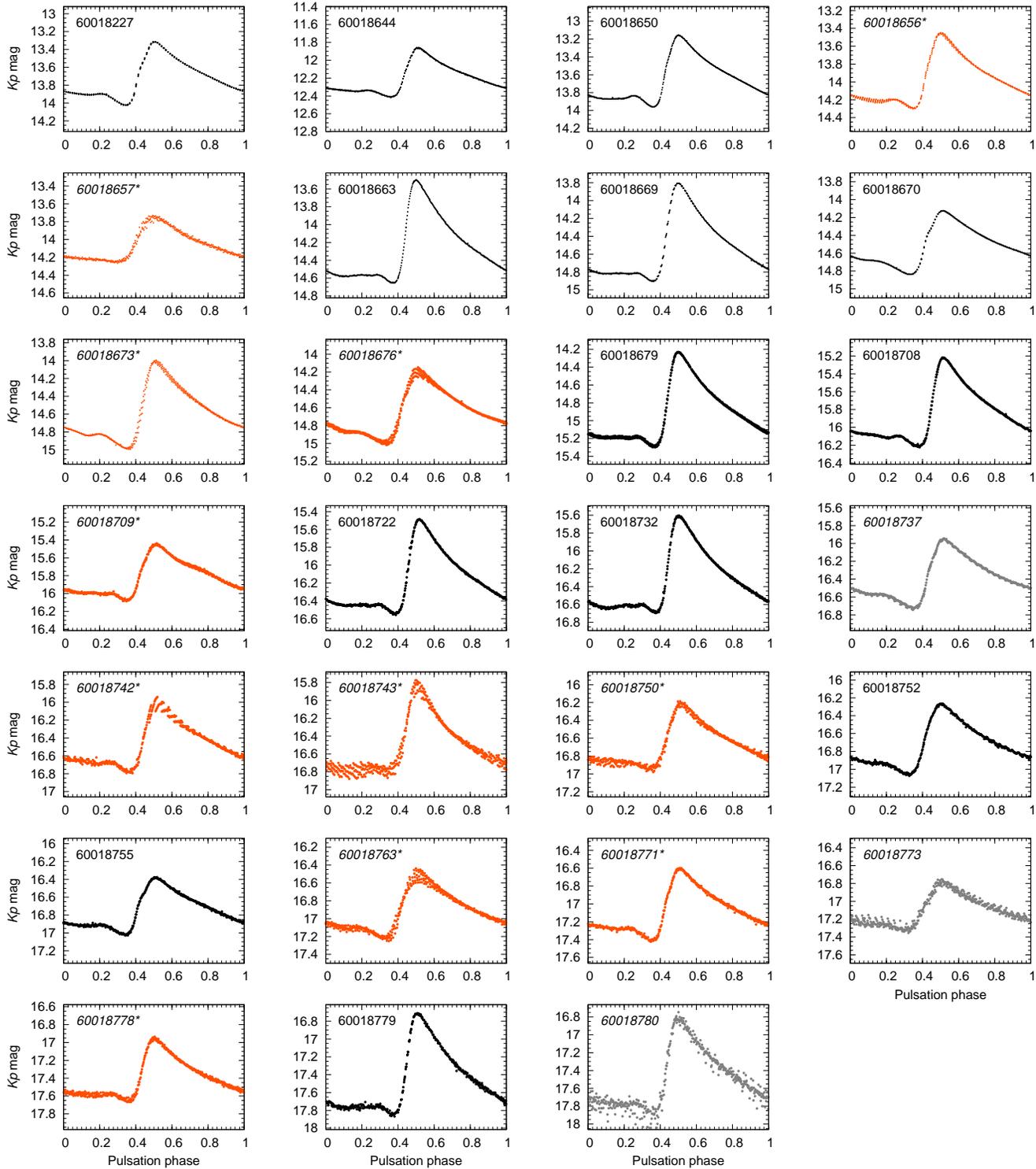


Figure 8. Folded light curves of the 27 RRab stars in the K2-E2 sample, based on the periods presented in Table 5. In black, roman labels: non-modulated stars; in orange, and marked with asterisks: Blazhko stars; in grey, italic labels: uncertain.

sparse sampling in this area with 2-300 data points scattered over 7-8 years. Each field is covered 3-4 times per night and fields repeat after one or two weeks. This kind of sampling is well suited for asteroid hunting but less useful to determine the Blazhko periods for RR Lyrae stars. The precision of individual CSS data points is 0.05 mag for bright targets

(~ 13 mag) but reaches 0.1 magnitudes for targets fainter than 17 mag that is well below the (sub)mmag accuracy of the K2 data.

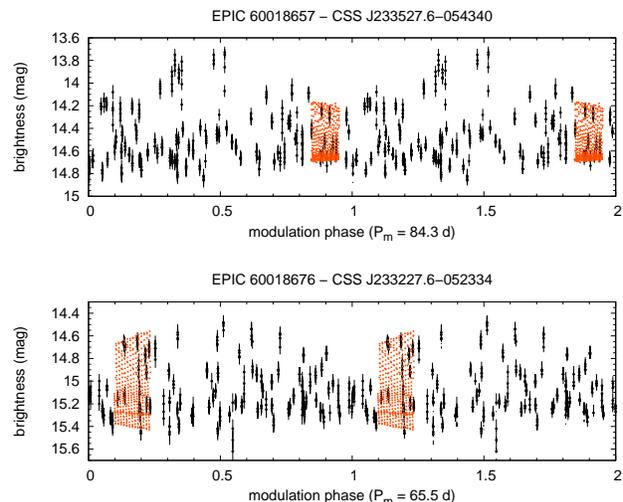
We searched for possible modulation sidepeaks in the Fourier spectra of the CSS data. In some cases the beating between the pulsation period and the sampling resulted in

Table 5. Fundamental-mode RR Lyrae stars observed during the Two-Wheel Concept Engineering Test run of the K2 mission. Photometric [Fe/H] indices are included as well. The average uncertainty of the indices is ± 0.1 dex.

EPIC ID	RA (deg)	Dec (deg)	K_p (mag)	Period (d)	Name	[Fe/H]	CSS ID No.	P_m (d)	Addtl. modes
(a) Non-modulated RRab stars									
60018227	0.83750	+3.39833	13.720	0.579041	ASAS J000321+0323.9	-1.30	—	—	—
60018644	354.15513	-2.21173	12.199	0.645043	ASAS J233637-0212.7	-1.88	1001126012320	—	<i>g</i> -mode?
60018650	356.29939	-5.20608	13.636	0.630289	NSVS 11906749	-1.95	1004127008648	—	—
60018663	355.36115	-5.54157	14.221	0.468800	NSVS 11904882	-1.34	1004127002457	—	—
60018669	351.61480	-2.97015	14.490	0.465721	CSS J232627.6-025814	-1.25	1004126052487	—	—
60018670	357.53961	+1.02409	14.501	0.669217	NSV 26151	-1.29	1101128018965	—	—
60018679	1.1001	-6.85657	14.888	0.494723	CSS J000424.0-065123	-1.30	1007001029423	—	—
60018708	356.48300	-7.47260	15.813	0.590650	CSS J234555.9-072821	-1.63	1007126018552	—	—
60018722	359.57559	-8.71836	16.133	0.482625	CSS J235818.1-084306	-1.29	1009127050514	—	—
60018732	352.76588	+1.06081	16.292	0.498292	CSS J233103.7+010339	-1.56	1101126021705	—	—
60018752	352.46439	+1.65528	16.710	0.594034	CSS J232951.4+013919	-1.26	1101126033293	—	(PD?)
60018755	1.94723	+3.26733	16.740	0.681966	CSS J000747.3+031602	-1.78	1104001007772	—	—
60018779	351.77498	-2.31793	17.415	0.465741	CSS J232705.9-021904	-1.33	1001126009779	—	—
(b) Modulated RRab stars									
60018656	355.98041	-4.05090	13.956	0.565113	CSS J234355.2-040303	-1.54	1004127030869	—	—
60018657	353.86545	-5.72775	14.054	0.543388	CSS J233527.6-054340	-1.38	1007125047274	84.3	PD, F2
60018673	357.75112	+4.87943	14.561	0.620309	CSS J235100.2+045245	-1.49	1104128039542	—	F2
60018676	353.11519	-5.39302	14.633	0.699815	CSS J233227.6-052334	-1.68	1004126005022	65.5	—
60018709	352.8671	-2.62876	15.815	0.563396	CSS J233128.1-023743	-1.19	1001126004067	—	—
60018742	1.90229	-8.24558	16.460	0.594456	CSS J000736.5-081444	-1.50	1007001004138	—	—
60018743	352.99328	-4.92403	16.481	0.493351	CSS J233158.3-045526	-1.75	1004126013659	—	PD
60018750	356.47678	-7.05360	16.658	0.548491	CSS J234554.4-070312	-1.36	1007126026371	—	—
60018763	357.49761	+4.62398	16.900	0.586196	CSS J234959.4+043726	-1.19	1104128034480	—	—
60018771	357.55888	+4.24208	17.065	0.553989	CSS J235014.1+041431	-1.22	1104128027645	—	—
60018778	355.94061	+1.86808	17.376	0.620342	CSS J234345.7+015205	-2.00	1101127037513	—	—
(c) Possibly modulated RRab stars									
60018737	355.31566	-4.21736	16.376	0.635342	CSS J234115.7-041302	-1.26	1004127027477	—	—
60018773	359.53266	+3.27389	17.094	0.595290	CSS J235807.8+031626	-1.30	1104128008817	—	—
60018780	1.56091	-4.50438	17.462	0.505453	CSS J000614.6-043015	-1.08	1004001019454	—	—

modulation-like variations. To exclude them, we compared the modulation phases between the CSS and K2 light curves. We were able to determine the Blazhko periods for two stars, EPIC 60018657 ($P_m = 84.3 \pm 0.3$ d) and 60018676 ($P_m = 65.5 \pm 0.3$ d). Despite the poor coverage, the final light curves in Figure 9 line up quite well. Interestingly, this method failed for two stars that show strong amplitude variations, EPIC 60018742 and 60018743. We suspect that these stars have shorter modulation periods than the two we presented above and the CSS observations are simply too sparse in time to identify those periods properly.

Finally, EPIC 60018709 shows variations in its light curve shape and pulsation amplitude. After prewhitening with the pulsation frequency and its harmonics, however, very little evidence can be found for side peaks anywhere but the $3f_1$ frequency value. Examination of the CSS data did not reveal any significant modulation side peaks, but the folded light curve indicates some amplitude variation around maximum light. Therefore we classified this star as a modulated RRab variable.

**Figure 9.** Folded Catalina Sky Survey and K2-E2 light curves of two modulated RRab stars where we were able to estimate the Blazhko periods.

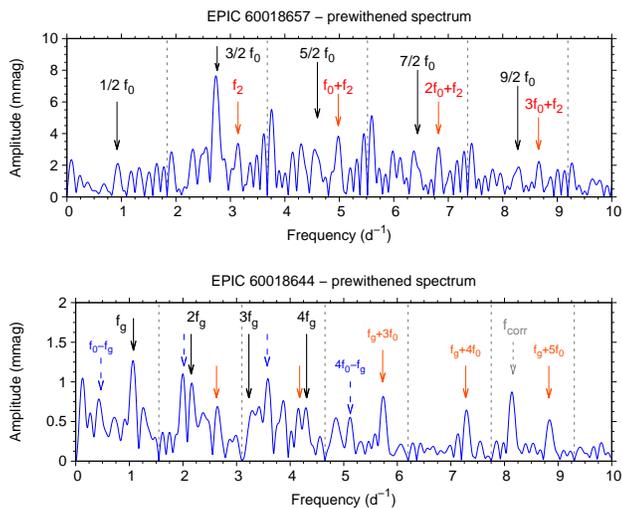


Figure 10. Top: residual Fourier spectrum of the modulated star EPIC 60018657 after prewhitening with the pulsation frequency and its harmonics (grey dashed lines). Identified or suspected additional peaks are labelled. In this case the Blazhko period is long (84.3 d), so the star shows little amplitude and phase variation over the K2 observations leading to very small modulation sidepeaks in the spectrum. Bottom: residual spectrum of EPIC 60018644, revealing the peaks corresponding to the g -mode suspected in the star. Thin solid and dashed arrows indicate the combination frequencies, $f_0 + nf_g$ and $nf_0 - f_g$, respectively. Some labels are not shown in the figure to avoid clutter.

4.1.2 Additional modes in RRab stars

One of the major discoveries of *Kepler*, regarding fundamental-mode RR Lyrae stars, was the identification of low-amplitude additional modes (Benkő et al. 2010, 2014). With the help of model calculations, we were able to trace back the origin of period doubling to one of these modes, the ninth overtone (Szabó et al. 2010; Kolláth et al. 2011). We note in passing that although the ninth overtone itself has very low amplitude and is undetectable in the K2-E2 sample, period doubling itself can be still observed since it is not a classical beating between modes but a purely dynamical phenomenon. The strong resonance itself destabilizes the limit cycle of the fundamental mode that bifurcates into a new, period-doubled limit cycle that takes two pulsation periods to repeat. This in turn creates a series of half-integer peaks in the Fourier spectrum but those do not correspond to actual pulsation modes themselves.

We searched thoroughly for additional modes in the K2-E2 data set. For classical variables, the widely used criterion of accepting a peak is $S/N > 4$. In the long data sets available from *CoRoT* and the *Kepler* prime mission, however, combinations with f_0 were always observed for the detected additional peaks. Therefore, we also considered a peak (or a series of peaks) reliable and of stellar origin when it formed combination peaks with the fundamental mode, e.g., a regular pattern was detectable in the spectrum even if the combinations had a S/N ratio between 2 and 4. Four authors searched for periods independently and the results were compared to each other: we only accepted peaks that were detected by at least two authors. The results are included in Table 5 and we detail them below.

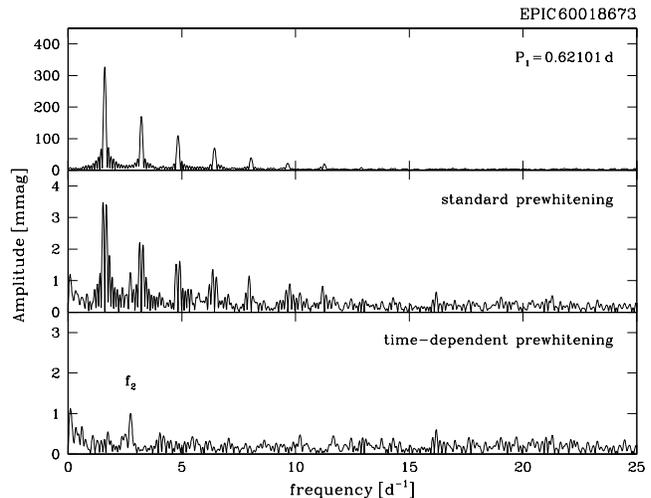


Figure 11. Fourier spectrum of the RRab star EPIC 60018673 (top), and the residual spectra. Middle: residual of the standard prewhitening with the main frequency and its harmonics: the strong, unresolved residual power hinders the detection of the additional modes. Bottom: residual of the time-dependent prewhitening. The f_2 peak is unambiguously detected.

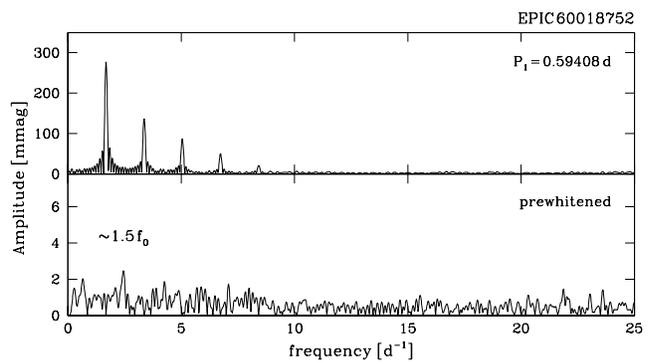


Figure 12. Fourier spectrum of the RRab star EPIC 60018752 (top), and the residual spectrum (bottom). A single half-integer peak ($1.5f_0$) is tentatively detected, suggesting that period doubling might occur in the pulsation.

We identified the same instrumental frequency peak that we found in EPIC 60018238, $f_{\text{corr}} \approx 8.16 \text{ d}^{-1}$, in two RRab stars, EPIC 60018227 and 60018644, and excluded it from the analysis.

Despite our efforts, we found period doubling in only two stars, EPIC 60018657 and 60018743, and we have tentative detection in one more faint target (see Figure 12). We now know, however, that the amplitude of period doubling is very variable, often on short time scales. Long-term observations provided by *CoRoT* and *Kepler* can help the detection by accumulating multiple high-amplitude phases and lowering the noise in the Fourier-spectra (Benkő et al. 2014; Szabó et al. 2014). On the other hand, short-term measurements, such as the K2-E2 can miss these signals entirely.

The Fourier spectrum of EPIC 60018657 revealed not only half-integer peaks ($(2n+1)/2 f_0$), the tell-tale signs of period doubling, but other additional peaks as well. As the residual spectrum in Figure 10 illustrates, the $f_2 = 3.14198 \text{ d}^{-1}$ mode and its linear combinations with the nf_0 frequen-

cies can be easily detected. Its frequency ratio, $f_0/f_2 = 0.5852$, is very similar to the ratios observed in other Blazhko stars in the *Kepler* and *CoRoT* samples (Benkő et al. 2014; Szabó et al. 2014). This mode is hypothesized to be the second radial overtone, although it was not reproduced by non-linear hydrodynamic models so far. The reason of its absence could be that it is in fact a non-radial mode close to the frequency of the second overtone. It can also be hypothesized that the excited mode exists in a parameter space that 1D models have avoided so far, similar to the discovery of period doubling and the low-amplitude first overtone in some R Rab models (Molnár et al. 2012). We note, however, that the mode was not detected during the extensive survey of those models. Some excess amplitude still remains in the spectrum after the removal of the half-integer and f_2 peaks, but the dataset is too short to resolve it into further additional modes. We detected the f_2 peak firmly in one other star, EPIC 60018673, at $f_2 = 2.7380 \text{ d}^{-1}$, with a period ratio of $f_0/f_2 = 0.5881$. This star shows stronger side peaks but we were able to detect f_2 with the time-dependent Fourier method, despite its small amplitude (1.1 mmag, Figure 11).

Additionally, we identified an intriguing periodicity in EPIC 60018644. Although the peak at $f = 1.0774 \text{ d}^{-1}$ itself has a signal-to-noise ratio of 3, we identified multiple significant peaks related to it (harmonics and combinations with f_0), therefore we accept this as a potential detection of a separate mode in the star. What makes this mode particularly interesting is that the period is longer than the fundamental mode itself, with a $P/P_0 = 0.695$ ratio, so it must be a g -mode that are always non-radial. Similar modes have been identified recently in first-overtone stars observed by *Kepler* and OGLE. The additional mode found in multiple OGLE stars has a period ratio of $\sim 0.92 - 0.95$ (scaled to the expected period of the fundamental mode) that is very different from our value (Netzel, Smolec & Dziembowski 2015). On the other hand, *Kepler* RRc stars show a multitude of modes between scaled period ratios from 0.38 to 0.97 (Moskalik et al. 2015).

As the last column of Table 5 indicates, we did not detect unambiguous additional p -modes in the non-modulated R Rab stars, in accordance with previous results (Benkő et al. 2014; Szabó et al. 2014), although a potential g -mode was identified in one star. Given the shortness of the data sets, however, both the tentative mode detections and the non-modulated status of the stars can be erroneous.

4.2 Photometric [Fe/H] indices

Light curve parameters, especially the ϕ_{31} Fourier phase relations of R Rab stars ($\phi_{31} = \phi_3 - 3\phi_1$, as defined by Simon & Lee 1981) can be used to calculate photometric metallicity indices. Nemeč et al. (2013) compared the Fourier parameters of the original *Kepler* RR Lyrae sample with indices derived from spectra obtained with the CFHT and Keck telescopes. We applied the same relation (Eq. 3 from Nemeč et al. 2013) to calculate the photometric [Fe/H] indices of the K2-E2 stars. The results are plotted in Figure 13 and included in Table 5. The K2-E2 sample is moderately metal-poor, with [Fe/H] values ranging from -1.08 to -2.00 . Based on the experiences of Nemeč et al. (2013), we estimate the accuracy of the indices to be ± 0.1 dex for non-modulated stars.

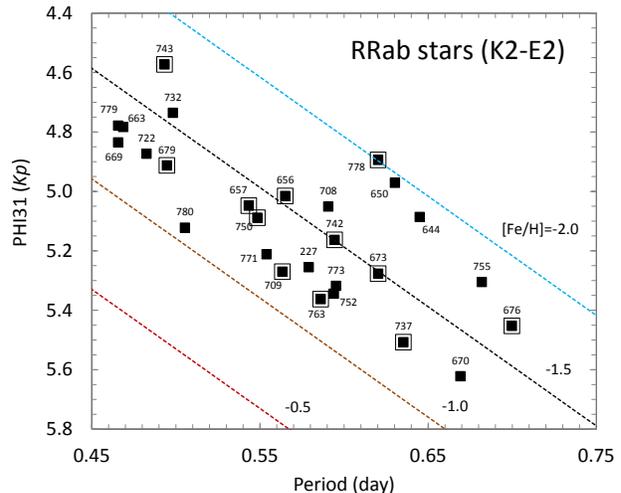


Figure 13. Fundamental periods versus ϕ_{31} Fourier phase relations of the R Rab stars. The relation of the two measures provide the photometric metallicities of the fundamental-mode stars that are included in Table 5. Modulated stars are indicated with additional boxes around the squares. Dashed lines correspond to various iso-metallicity lines, based on Eq. 3 of Nemeč et al. (2013).

The accuracy is likely more limited for the modulated stars where the phase relation values change over the Blazhko cycle, but it is still acceptable. Smolec (2005) showed that filtering out the modulation side-peaks can lead to reasonably good [Fe/H] values. Nemeč et al. (2013) also investigated the relation between the spectroscopic and photometric [Fe/H] determinations and found that the values are consistent for most Blazhko stars. Only three extremely modulated stars, less than 10 per cent of the overall sample, showed significant discrepancies. Based on the original *Kepler* sample, we can estimate that about 90 per cent of the K2-E2 stars should closely follow the relation we used.

We consider the limited modulation phase coverage the strongest source of uncertainty in our study. Both the pulsation period and ϕ_{31} change during the modulation cycle: based on the values presented by Nemeč et al. (2013), we estimate the overall uncertainty of the [Fe/H] values for the Blazhko stars to be $\pm 0.15 - 0.2$ dex.

5 CONCLUSIONS AND FUTURE PROSPECTS

The *Kepler* space telescope observed 1952 targets in Pisces for 8.9 days during the K2 Two-Wheel Concept Engineering Test run. We identified 32 RR Lyrae stars among the targets and one more as a background object in the target pixel mask of another star. We extracted the light curves from the target pixel files with the PYKE software tool (Still & Barclay 2012). The analysis of the photometric data yielded the following results:

- We identified 2 double-mode (RRd), 4 first-overtone (RRc) and 27 fundamental-mode (RRab) stars in the field. Most of the stars were discovered by the Catalina Sky Survey (Drake et al. 2014).
- The f_X additional mode was detected in the two RRd and the three non-modulated RRc stars, although with vary-

ing degrees of significance. The mode appears to be period-doubled and varies over time in three out of five stars.

- One of the RRc stars, CSS J125742.1–015022, turned out to be a modulated first-overtone star. This is the first time that a Blazhko-RRc star was observed by a space photometric mission. The modulation period of the star is 17.07 d and both the pulsation amplitude and phase show strong variations. It is the only RRc star where no additional modes were found.

- Indications of a potential g -mode was found in one RRab star, EPIC 60018644, at a period ratio of $P_0/P_g = 0.695$.

- Even from such short dataset, we found indications that at least 11 out of 27 RRab stars (41 per cent) may be modulated. We detected period doubling and the (presumed) second overtone in multiple stars. Based on the Fourier-parameter relations, the $[\text{Fe}/\text{H}]$ values of the stars fall between -1.08 and -2.00 dex.

- As a by-product of the target search for the K2 campaigns, we re-examined the ‘RRd’ and ‘Blazhko’ classifications of the Catalina Periodic Variable Star Catalog and identified 165 bona-fide double-mode stars throughout the sky, 130 of which were previously unknown. For a brief description, see Appendix A.

The results promise a bright future for the K2 mission. Even from these very short datasets, we were able to identify all new phenomena that space-based photometric missions discovered in RR Lyrae stars so far. The 75-day long campaign data will provide much better frequency resolution and refined light curve extraction methods will improve the photometric precision.

Losing the ability of *Kepler* to observe stars for several years was unfortunate and will be very hard to reproduce. The silver lining of K2 mission, however, is its flexibility to observe many different stars. This single field contained 33 RR Lyrae variables. Based on our estimates and target selection so far, this number of expected RR Lyrae stars will vary between approximately 10 and 200 for the forthcoming fields (Molnár, Plachy & Szabó 2014). With a sustained observing program, we will be able to answer new questions the previous missions could not: what is the true ratio of modulated RRc stars? Are all non-modulated RRab stars strictly periodic? Is the f_x mode present in all RRd and RRc stars? Finally, K2 will be able to detect RR Lyrae stars in globular clusters, halo streams, and hopefully in nearby dwarf galaxies as well, broadening the scope of the mission even further.

ACKNOWLEDGMENTS

This paper includes data collected by the *Kepler* spacecraft during the K2 ‘Two-wheel Concept Engineering Test’ operations. The authors gratefully acknowledge the *Kepler* team, the Guest Observer Office, and Ball Aerospace, whose outstanding efforts have made these results possible. Funding for the *Kepler* spacecraft is provided by the NASA Science Mission Directorate. This work has been supported by the ESA PECS Contract No. 4000110889/14/NL/NDe, and the Lendület-2009 and LP2014-17 Young Researcher Programs of the Hungarian Academy of Sciences. The research leading to these results has received funding from the European

Community’s Seventh Framework Programme (FP7/2007-2013) under grant agreements no. 269194 (IRSES/ASK), no. 312844 (SPACEINN), and no. 338251 (StellarAges). PM and RS are supported by the Polish National Science Center through grant DEC-2012/05/B/ST9/03932. CRTS and CSDR2 are supported by the U.S. National Science Foundation under grant AST-1313422. KK is grateful for the support of Marie Curie IOF grant 255267 SAS-RRL (FP7). The CSS survey is funded by the National Aeronautics and Space Administration under Grant No. NNG05GF22G issued through the Science Mission Directorate Near-Earth Objects Observations Program. This research has made use of the SIMBAD database, operated at CDS, Strasbourg, France.

REFERENCES

- Benkő J. M., et al., 2010, MNRAS, 409, 1585
 Benkő J. M., Plachy E., Szabó R., Molnár L., Kolláth Z., 2014, ApJS, 213, 31
 Bernard K., Wils P., 2006, IBVS, 5698, 1
 Borucki W. J. et al., 2010, Science, 327, 977
 Brown T. M., Ferguson H. C., Smith E., Kimble R. A., Sweigart A. V., Renzini A., Rich R. M., 2004, AJ, 127, 2738
 Buchler J. R., Kolláth Z. 2011, ApJ, 731, 24
 Chadid M., et al., 2010, A&A, 510, 39
 Chadid M., 2012, A&A, 540, 68
 Clementini, G., et al. 2000, AJ, 120, 2054
 Conroy K. E., et al., 2014, PASP, 126, 914
 Cseresnjes, P., 2001, A&A, 375, 909
 Derekas A., Kiss L. L., Udalski A., Bedding T. R., Szatmáry K., 2004, MNRAS, 354, 821
 Drake A. J., et al. 2009, ApJ, 696, 870
 Drake A. J., et al. 2013a, ApJ, 763, 32
 Drake A. J., et al. 2013b, ApJ, 765, 154
 Drake A. J., et al. 2014, ApJS, 213, 9
 Dziembowski, W. A., 2012, A&A, 62, 323
 Guggenberger E., et al., 2012, MNRAS, 424, 649
 Gruberbauer M. et al., 2007, MNRAS, 379, 1498
 Hajdu G., Catelan, M., Jurcsik, J., Dkny, I., Drake, A. J., Marquette, J.-B., 2015, MNRAS, 449, L113
 Hamsch F.-J., Wils P., 2008, IBVS, 5840, 1
 Haschke R., Grebel E. K., Duffau S., 2012, AJ, 144, 106
 Hippke M., Learned J. G., Zee A., Edmondson W. H., Steven I. R., Lindner J. F., Kia B., Ditto W. L., 2015, ApJ, 798, 42
 Howell S. B., et al., 2014, PASP, 126, 398
 Jurcsik J., Sódor Á., Szeidl B. et al., 2009, MNRAS, 400, 1006
 Jurcsik J., Smitola P., Hajdu G., et al., 2015, ApJS, accepted, arXiv:1504.06215
 Kolenberg K. et al., 2010, ApJ, 713, 198
 Kolenberg K. et al., 2011, MNRAS, 411, 878
 Kolláth Z., Buchler J. R., 2001, in *Stellar Pulsation – Non-linear Studies*, Astrophysics and Space Science Library Series vol. 257, pp. 29-60, Kluwer Academic Publishers, Dordrecht
 Kolláth Z., Molnár L., Szabó R., 2011, MNRAS, 414, 1111
 Kolláth Z., Buchler J. R., Szabó R., Csabry Z., 2002, A&A, 385, 932

Le Borgne J. F., et al., 2014, MNRAS, 441, 1435
 Lenz P., Breger M., 2005, CoAst, 146, 53
 Mizerski T., 2003, Acta Astr., 53, 307
 Molnár L., Kolláth Z., Szabó R., Plachy E., 2012, AN, 333, 950
 Molnár L., Plachy E., Szabó R., 2014, IBVS, 6108, 1
 Moskalik P., Kołaczowski Z., 2008, MNRAS, 394, 1649
 Moskalik P. et al., 2013, ASSP, 31, Supp. Mat. P34, arXiv:1208.4251
 Moskalik P. et al., 2015, MNRAS, 447, 2348
 Nagy A., Kovács G., 2006, A&A, 454, 257
 Nemeč J. M. et al., 2011, MNRAS, 417, 1022
 Nemeč J. M., Cohen J. G., Ripepi V., Derekas A., Moskalik P., Sesar B., Chadid M., Bruntt H., 2013, ApJ, 773, 181
 Netzel H., Smolec R., Dziembowski W., 2015, MNRAS, 451, L25
 Netzel H., Smolec R., Moskalik P., 2015, MNRAS, 447, 1153
 Plachy E., Benkó, J. M., Kolláth Z., Molnár L., Szabó R., 2014, MNRAS, 445, 2810
 Poleski R., 2014, PASP, 126, 509
 Sarajedini A., Barker M. K., Geisler D., Harding P., Schommer R., 2006, AJ, 132, 1361
 Sesar B. et al., 2013, AJ, 146, 21
 Simon N. R., Lee A. S., 1981, ApJ, 248, 291
 Skarka M., 2013, A&A, 549, 101
 Skarka M., Zejda M., 2013, MNRAS, 428, 1442
 Smolec, R., 2005, AcA, 55, 59
 Smolec R., Moskalik P., 2008, AcA, 58, 233
 Smolec R. et al., 2015, MNRAS, 447, 3756
 Soszyński, I., et al., 2011, AcA, 61, 1
 Soszyński, I., et al., 2014, AcA, 64, 177
 Still M., Barclay T., 2012, Astrophysics Source Code Library, ascl:1208.004
 Szabó R., Kolláth Z, Buchler J. R., 2004, A&A, 425, 627
 Szabó R. et al. 2010, MNRAS, 409, 1244
 Szabó R. et al. 2014, A&A, 570, 100
 Szczygiel D. M., Fabrycky D. C., 2007, MNRAS, 377, 1263
 Van Cleve J., Caldwell D., 2009, *Kepler Instrument Handbook*, KSCI 19033-001 (Moffett Field, CA: NASA Ames Research Center)
 Vanderburg A., Johnson J. A., 2014, PASP, 126, 948
 Wils P., 2010, IBVS, 5955, 1
 Wils P., Otero, S. A., 2005, IBVS, 5593, 1
 Wils P., Kleidis S., Broens E., 2008, MNRAS, 387, 783
 Wils P., Lloyd C., Bernhard K., 2006, MNRAS, 368, 1757

APPENDIX A: DOUBLE-MODE STARS IN THE CATALINA SKY SURVEY

As a side-product to the K2 investigations we also identified 130 new RRd stars throughout the sky. These are very valuable targets, because their physical parameters can be estimated from hydrodynamic model calculations, as section 3.1.1 illustrates. The Catalina Sky Survey provided us perhaps with the richest database of RR Lyrae stars for target selection for all K2 fields that are not too close to the Galactic plane. As usually only 1-5 RRd stars can be observed in a K2 campaign, proper identification is especially important. This work is also the continuation of the search for bona-fide RRd stars in the LINEAR sample by Poleski

(2014). The stars we found are distributed between declinations -30° and 60° , except for a ~ 30 -degree band along the Galactic plane (Figure A1).

Double-mode RR Lyr stars in the Catalina sample were separated from other RR Lyr by Drake et al. (2014), based on the visual examination of the light curves. No prewhitening was performed, thus, such classification can be uncertain as indicated by the misclassification of the two stars in the K2-E2 field. As the rare RRd stars are valuable targets, we searched for bona-fide double-mode pulsators among the stars that Drake et al. (2014) classified as ‘RRd’ or ‘Blazhko’ stars. The photometric data were downloaded from <http://catalinadata.org/> (Drake et al. 2009). We found 461 stars classified as RRd and 70 stars classified as Blazhko for which at least 100 data points were obtained. Double mode pulsations were found in 161 and 5 stars, respectively. This means that only 35 per cent of the stars classified as ‘RRd’ in Catalina catalogue are in fact double-mode pulsators. We checked that the fraction of confirmed RRd stars does not change if we limit to stars with at least 300 data points. Most of the stars that were erroneously identified as ‘RRd’ in the CSS variable catalogue actually show Blazhko-type modulation. One star turned out to have two different CSS identifications (CSS_J104003.0+414504 = CSS_J104003.2+414503), bringing the total number of confirmed RRd stars to 165. We checked the literature for these stars and found that 35 of them were previously identified, most in the LINEAR survey data (29 stars, Poleski 2014). Six stars were previously found by Clementini et al. (2000); Bernard & Wils (2006); Szczygiel & Fabrycky (2007); Hamsch & Wils (2008), and Wils (2010). The number of confirmed RRd stars is increased by 130. Table A1 presents all confirmed RRd stars and Table A2 lists their pulsation parameters. Figure 2 compares the positions of newly identified stars on Petersen diagram with RRd stars from various other sources.

We also checked CSS photometry for the two K2-E2 RRd stars. After removing outlying data points from Catalina data we were able to confirm double-mode pulsations in both stars. This illustrates that more double-mode stars are hidden in Catalina lists of the RRab and RRc stars. However, those lists are much longer, with tens of thousands and more than 5000 stars, respectively, and will require further work to classify.

Table A1. Sample table of the confirmed double-mode RR Lyrae stars in the Catalina Surveys Periodic Variable Stars Catalog. The list includes 165 stars, based on the examination of the light curves of the stars originally labelled as ‘RRd’ or ‘Blazhko’ in the Catalog, along with the two K2 stars. Full table is available in the online version.

CSS name	RA (deg)	Dec (deg)	Brightness (mag)	CSS ID	Other ID
CSS_J001420.8+031214	3.586833	+3.20389	17.45	1104002007409	—
CSS_J001724.9+200542	4.353750	+20.09506	16.64	1121002007726	—
CSS_J001812.9+210201	4.554042	+21.03375	14.54	1121002027610	—
CSS_J001836.2-191522	4.651083	-19.25619	16.84	1018002009928	—
CSS_J003359.4+022609	8.497833	+2.43583	15.87	1101004049971	—
CSS_J004804.2+365310	12.017833	+36.88625	16.60	1138004006108	—
CSS_J011032.5+215708	17.635500	+21.95233	16.13	1121006050243	—
CSS_J011824.3-123318	19.601583	-12.55525	17.72	1012007028731	—
CSS_J014305.3+010549	25.772167	+1.09708	17.04	1101010021310	—
...					
CSS_J082622.9+282405	126.595625	+28.40142	16.84	1129040004714	LINEAR 658512
...					

Table A2. Sample table of the pulsation parameters of the confirmed double-mode RR Lyrae stars. FM: fundamental mode, O1: first overtone. Full table is available in the online version.

CSS name	FM period (day)	FM period uncert.	FM ampl. (mag)	O1 period (day)	O1 period uncert.	O1 ampl. (mag)	Per. ratio (P_1/P_0)
CSS_J001420.8+031214	0.51930994	0.00000166	0.23	0.38711458	0.00000056	0.39	0.74544034
CSS_J001724.9+200542	0.48002554	0.00000174	0.29	0.35712249	0.00000074	0.41	0.74396560
CSS_J001812.9+210201	0.55837503	0.00000198	0.19	0.41615750	0.00000064	0.34	0.74530106
CSS_J001836.2-191522	0.46299846	0.00000198	0.28	0.34403253	0.00000067	0.40	0.74305329
CSS_J003359.4+022609	0.48368923	0.00000208	0.21	0.36009918	0.00000087	0.29	0.74448459
CSS_J004804.2+365310	0.46874750	0.00000070	0.38	0.34831989	0.00000079	0.37	0.74308640
CSS_J011032.5+215708	0.47302180	0.00000138	0.33	0.35178257	0.00000051	0.40	0.74369209
CSS_J011824.3-123318	0.52620846	0.00000320	0.27	0.39247062	0.00000147	0.38	0.74584628
CSS_J014305.3+010549	0.47555267	0.00000161	0.32	0.35372017	0.00000058	0.40	0.74380861
...							
CSS_J082622.9+282405	0.47390930	0.00000138	0.22	0.35260694	0.00000044	0.35	0.74403887
...							

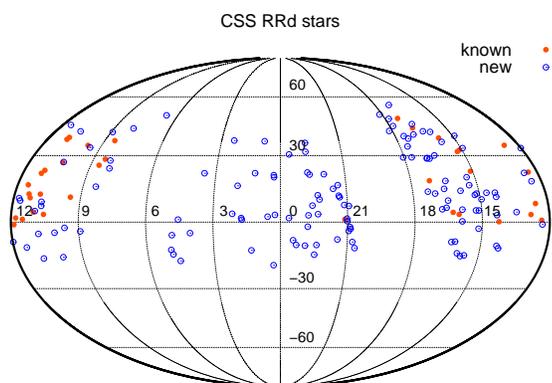


Figure A1. Distribution of the CSS double-mode RR Lyrae stars in the sky, in Equatorial coordinates. Blue circles are the new discoveries, orange dots are the already known stars.